

On the multispacecraft determination of periodic surface wave phase speeds and wavelengths

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Received 13 December 2009; revised 9 February 2010; accepted 8 March 2010; published 2 September 2010.

[1] Observations of surface waves on the magnetopause indicate a wide range of phase velocities and wavelengths. Their multispacecraft analysis allows a more precise determination of wave characteristics than ever before and reveal shortcomings of approximations to the phase speed that take a predetermined fraction of the magnetosheath speed or the average flow velocity in the boundary layer. We show that time lags between two or more spacecraft can give a qualitative upper estimate, and we confirm the unreliability of flow approximations often used by analyzing a few cases. Using two-point distant magnetic field observations and spectral analysis of the tailward magnetic field component, we propose an alternative method to estimate the wavelength and phase speed at a single spacecraft from a statistical fit to the data at the other site.

Citation: Foullon, C., C. J. Farrugia, A. N. Fazakerley, C. J. Owen, F. T. Gratton, and R. B. Torbert (2010), On the multispacecraft determination of periodic surface wave phase speeds and wavelengths, *J. Geophys. Res.*, *115*, A09203, doi:10.1029/2009JA015189.

1. Introduction

[2] The frequent motions of the magnetopause and of its adjacent boundary layers can be of different origins. Clearly, they may be caused by a variable shocked solar wind plasma impinging on this region. However, they may also be due to inherent instabilities at the boundary itself, brought about for example by flow shears, in the case of the Kelvin-Helmholtz (KH) instability [e.g., *Southwood*, 1968; *Hasegawa*, 1975; *Farrugia et al.*, 1998]. The latter mechanism, in particular, may lead in its nonlinear stage to plasma entry to the magnetosphere, by entraining oppositely directed magnetic field lines in the vortical structures, which it gives rise to [e.g., *Hasegawa et al.*, 2004]. Although it has been investigated on multiple levels of time-dependent simulations [e.g., *Miura*, 1984; *Wu*, 1986; *Manuel and Samson*, 1993; *Thomas and Winske*, 1993; *Fujimoto and Terasawa*, 1994], this mechanism needs a violation of the frozen-in law or magnetic reconnection triggered in KH vortices [e.g., *Otto and Fairfield*, 2000; *Nykyri and Otto*, 2004]. Wave characteristics, such as the sizes of KH vortices, and other derived quantities may be of physical importance, notably via comparison with theoretical and numerical models, to

elucidate the conditions leading to the waves' formation, their possible non-linear development and their role in the plasma entry to the magnetosphere. For these reasons, reliable estimates of the wave characteristics are important.

[3] Quoted values of surface wave observations on the magnetopause indicate that their periods are typically in the Pc5 (1–10 mHz) range, with a wide range of phase velocities (varying from about 60 km s⁻¹ [see *Foullon et al.*, 2008] to 350 km s⁻¹ [see *Chen et al.*, 1993]) and a similar spread in wavelengths (from 2 R_E [see *Foullon et al.*, 2008] to a few tens of R_E [see *Hones et al.*, 1981]). Arguably, a key factor controlling the wavelengths, λ , is the distance of the observing site from the subsolar point because the magnetosheath flow speed picks up as one moves tailward, leading to λ -stretching effect, and the magnetic field strength decreases along the magnetopause, implying that all gyroradii increase. In addition, KH surface waves in the magnetospheric context can become nonlinear while propagating down the tail. Nonlinear effects have been invoked to account for (1) wavelengths of a few R_E typically observed on the magnetopause, which, as argued by *Belmont and Chanteur* [1989], are much longer than those predicted by linear theory; (2) the tailward steepening of the KH leading fronts observed with Cluster [*Owen et al.*, 2004; *Foullon et al.*, 2008], consistent with the growing phase of KH waves [*De Keyser et al.*, 2005] (note that the inverse dependence found between the boundary layer thickness and the tailward steepening of the leading edge [*Foullon et al.*, 2008] suggest that this effect is affected by changing conditions of the medium in which the waves propagate); (3) the presence of vortices, a phenomenon supported by the interpretation of data in single or multispacecraft analyses [*Hones et al.*, 1981; *Saunders et al.*, 1983; *Fairfield et al.*, 2000; *Hasegawa et al.*,

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2004; *Fairfield et al.*, 2007]; (4) an inverse dependence between the interplanetary magnetic field (IMF) clock angle and the wavelength at the flank [*Foullon et al.*, 2008] or the period of the magnetospheric micropulsations these waves excite [*Farrugia et al.*, 2000], which confirms the significance of source regions and nonlinear development for interpreting observations of remotely generated KH waves.

[4] In the simplest linear magnetohydrodynamic (MHD) description, the phase speed of a KH wave propagating with wave vector \mathbf{k} in an ideal incompressible plasma, with a discontinuous velocity shear layer and assuming the layer to be infinitely thin [*Chandrasekhar*, 1961; *Hasegawa*, 1975], is

$$v_p = \frac{n_1 \mathbf{k} \cdot \mathbf{V}_1 + n_2 \mathbf{k} \cdot \mathbf{V}_2}{k(n_1 + n_2)}. \quad (1)$$

Here the indices refer to the two plasma environments on either side of the boundary, n is the plasma number density, \mathbf{V} is the plasma flow velocity; \mathbf{k} and \mathbf{V} (and the magnetic field vector, \mathbf{B} , not present in equation (1)) are all tangential to the layer. Equation (1) is the real part, $v_p = \omega/k$, in the dispersion relation associated with complex roots, $\sigma = \omega + i\gamma$, for perturbations along the boundary surface, i.e., $\exp(i(\mathbf{k} \cdot \mathbf{r} - \sigma t))$. The imaginary part (not shown), which corresponds to the growth rate of the instability, γ , is a competition between the destabilizing effects of velocity gradient or shear, $|\mathbf{V}_1 - \mathbf{V}_2|$, and the stabilizing effects from magnetic tension forces; the KH instability criterion is met for $\gamma > 0$. Several observations show a set of \mathbf{k} directions of predominantly plane disturbances consistent with the conditions predicted by this ideal KH instability criterion [e.g., *Ogilvie and Fitzenreiter*, 1989; *Owen et al.*, 2004].

[5] While this treatment is only approximate and qualitative as applied to the magnetospheric boundary and to nonlinear (vortical) and multiwavelength structures, the dimensions of those structures may be better estimated with the group velocity, which is the energy propagation velocity of waves of different wavelengths and propagation directions. Considering the convectively unstable surface mode [e.g., *Mills et al.*, 2000; *Wright et al.*, 2000], perturbations along the boundary surface, while growing, are convected downstream with the phase velocity $v_p(\mathbf{k})$ and the group velocity, \mathbf{v}_g , which from equation (1) corresponds to

$$\mathbf{v}_g = \frac{\partial \omega}{\partial \mathbf{k}} = \frac{n_1 \mathbf{V}_1 + n_2 \mathbf{V}_2}{n_1 + n_2}. \quad (2)$$

The group velocity is the propagation velocity of the wave envelope. Thus if there is a case for a coherent, nondispersive structure or wave envelope propagating in a particular direction, then the speed, V_k , that connects its leading and trailing front (i.e., speed of the wave envelope) is as meaningful as the group speed for calculating the wavelength of that structure. In this work, such observational nondispersive ‘phase speeds’ and associated propagation directions are discussed.

[6] Besides the stream of flow velocity, other factors may control λ , including (1) the thickness of the boundary layer [e.g., *Walker*, 1981; *Miura and Pritchett*, 1982], (2) the nonlinear development of dominant structures [e.g., *Belmont and Chanteur*, 1989; *Miura*, 1999a, 1999b], and (3) the

global magnetospheric geometry of unstable source regions controlled by the IMF orientation [e.g., *Farrugia et al.*, 1998; *Foullon et al.*, 2008]. All those factors have therefore a direct effect on V_k .

[7] Despite those latest theoretical developments, most observational, as well as theoretical works, assume some ‘flow approximations’, which may be seen to correspond to limiting cases to equation (2). Focusing on the KH waves, we first outline the advantages and results of multispacecraft timing, using between 2 and 4 spacecraft. We then compare and discuss some flow approximations often used, where we note clear disagreements, and thus propose alternative consistency checks to improve the estimates.

2. Surface Wave Phase Speeds From Multispacecraft Observations

[8] Triangulation or four-spacecraft timing analysis [*Russell et al.*, 1983; *Schwartz*, 1998] has been applied to our knowledge in two Cluster studies of surface waves, that by *Owen et al.* [2004] and that by *Foullon et al.* [2008], on the dawn and dusk flanks respectively. The waves studied have a sawtooth shape, which they retain as they propagate past Cluster (within a relatively short time interval), at an epoch when the four satellites have separations of ~ 2000 km. Such conditions are very appropriate for estimating their phase speeds, and consequently their wavelengths, via the method proposed by *Foullon et al.* [2008] (see illustration by *Foullon et al.* [2009] and also a more approximative variant in the work of *Owen et al.* [2004]). This method uses aggregate results from four-spacecraft timing analysis applied to a pair of bounding surfaces.

[9] Figure 1 explains a 2-D version of the geometries considered. The wavefront is shown by the vertical dotted line and so the propagation direction \mathbf{k} is along the horizontal direction. The separation distance along \mathbf{k} between two samplings of the wavefront in a 1 s time interval corresponds therefore to the value of the phase speed, V_k (in km s^{-1}). For a surface (plain line) inclined with respect to the wavefront, the separation distance is the same (Figure 1a). If the two sampling spacecraft are not aligned along \mathbf{k} , then the distance between the two samplings of what appears to be the same inclined surface (with the same given phase) is covered in 1 s with a larger apparent speed (Figure 1b). This later configuration corresponds to speeds along normals of bounding surfaces, typically inclined with respect to the wavefront. In a 3-D geometry, such speeds and surface normals can be obtained via four-spacecraft timing and the speeds are still larger than the periodic surface wave phase speed. The two-point and four-point measurements have the same problem of the apparent speed. The periodic surface wave phase speed is obtained via ‘deprojection’ along the wave propagation direction.

[10] The former configuration (Figure 1a) corresponds to the case of two spacecraft aligned along the magnetospheric boundary and presumably along \mathbf{k} , with the phase speed obtained by taking the spacecraft separation distance divided by the time lag between the two observations, e.g., ISEE-1 and ISEE-2 [*Hones et al.*, 1981; *Chen et al.*, 1993]. Even with the scale size of four spacecraft (tetrahedral) configuration (more precisely the tetrahedron scale size along the average direction tangential to the boundary), it is

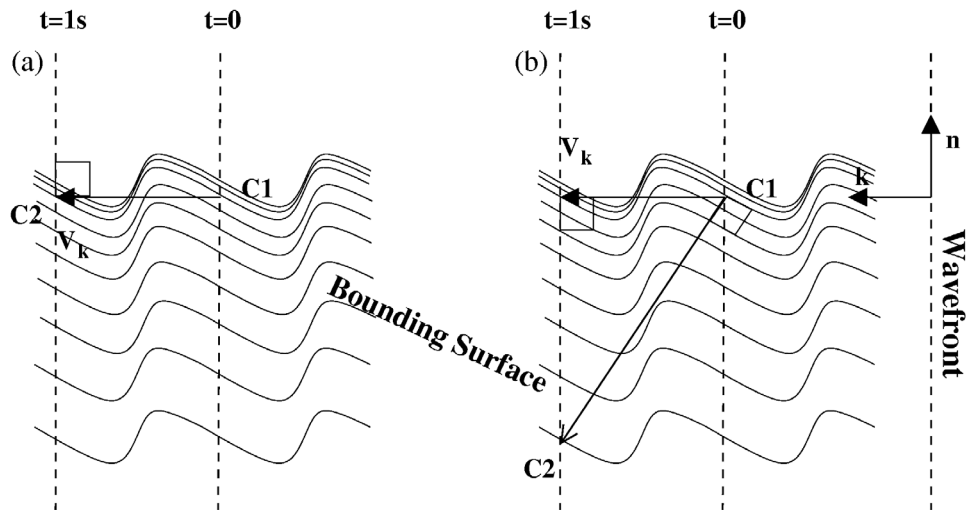


Figure 1. Surface wave propagating in the k direction with phase speed, V_k , showing bounding surfaces sampled by spacecraft C1 at time $t = 0$ and C2 at $t = 1$ s, respectively, (a) without and (b) with inclination of the two-spacecraft direction with respect to k .

possible to give an optimal estimate of the phase speed by considering the largest time lag between spacecraft. For instance, taking the ratio of the Cluster scale size of 2000 km over the largest time lag of 30 s [see *Foullon et al.*, 2008], yields a value of about 67 km s^{-1} , in good agreement with the more accurate results derived from four-spacecraft timing analysis. This can therefore be used to give a qualitative and immediate validation of the more comprehensive method. As illustrated in Figure 1b, the ‘in-between spacecraft’ speed may become an upper estimate to the real phase speed.

[11] Both Cluster studies of surface waves [*Owen et al.*, 2004; *Foullon et al.*, 2008] refer to similar distances from the subsolar point (with spacecraft located at $X_{GSM} \approx [-3, -6] R_E$, on opposite sides of noon and at low geomagnetic latitudes) and indicate phase speeds ($50\text{--}90 \text{ km s}^{-1}$) and wavelengths ($2\text{--}3.4 R_E$) in the lower range compared to other reports in the same locales. Accurate determination of the required time lags for a main front passage can be achieved via the use of both (1) plasma and (2) magnetic field data sets. First, plasma data sets that display sharp transitions between adjacent plasma environments have to be available on four spacecraft, e.g., the electron temperature from PEACE [*Owen et al.*, 2004; *Foullon et al.*, 2008]. Second, to deal with magnetic field data sets, which are more highly perturbed, *Foullon et al.* [2008] proposed a multiscale approach to select recurring local extrema, for different scales or levels of smoothing, in the time derivative of a magnetic field (sunward) component. The reliability of those timings is further cross-checked with the use of error estimates on the combined (plasma and magnetic field) results. For the dusk event, the errors on the phase speed range from 3 to 28 km s^{-1} . The phase speed from the Cluster results represents therefore a small fraction (a third or less) of the magnetosheath flow speed, V_s ($V_s \sim 260 \text{ km s}^{-1}$ in the work of *Foullon et al.* [2008]).

[12] In the case of surface waves, the timings obtained locate the same phase in the waveform but not necessarily

for the same isosurface (or discontinuity). The application of four-spacecraft timing analysis in this case thus differs from the case of shocks [e.g., *Foullon et al.*, 2007], where the propagation along the normal can be expected. Nevertheless, the technique can be applied as long as the plane surface hypothesis is valid, i.e., a quasi-planar subsection of the waveform is sampled, when the spacecraft separation is relatively small. This can be the case also in the presence of rolled-up vortices, for detected planar disturbances that embed the vortices. Finally, the effect of overall motions in response to varying conditions, which may interfere with the measurements, can be mitigated by taking the average of phase speeds and propagation directions between pairs of inward and outward motions [*Foullon et al.*, 2008]. This average can also result in an approximate direction of propagation, if the k direction changes notably across the boundary layer. The directions of propagation obtained in both Cluster studies are found to be perpendicular to an average external magnetic field direction, rather than being related to a magnetosheath flow component, suggesting that the wave propagation direction may adjust to the bending of field lines developed by the KH waves [*Foullon et al.*, 2010]. Comparisons with the single-spacecraft method introduced by *De Keyser and Roth* [2003] would be desirable. This latter method estimates the surface wave phase speed for which the flow in a frame comoving with the wave is parallel to the density gradient (usually in the boundary layer).

3. Flow Approximations Versus Estimates From Time Lags

[13] Another multispacecraft approach was proposed by *De Keyser et al.* [2004] [see also *De Keyser et al.*, 2005]. It uses an empirical reconstruction technique to obtain a 2D picture of the surface waves in a reference frame that slides along the boundary with the wave. Here the phase speed of the wave is assumed. For an event they examined on 25 November 2001, for instance, the phase speed is taken to

be equal to one half of the magnetosheath speed, i.e., $V_k = 200 \text{ km s}^{-1}$. However, the latter value does not concur with an (upper) estimate of about 100 km s^{-1} , which we infer from the 2500 km scale size of the spacecraft separation and the 25 s maximum time lags for the same event (seen in EFW densities [De Keyser *et al.*, 2005, Figure 8.37]).

[14] Many researchers stipulate the phase speed as a given fraction of V_s . This can lead to erroneous results and we are here trying to show how to improve these estimates. There are ‘anomalous’ cases of magnetopause waves with opposite (sunward) steepened edges reported from two-spacecraft observations by ISEE-1 and ISEE-2 [Chen *et al.*, 1993; Chen and Kivelson, 1993]. These anomalous waves appear to form in association with strong magnetosheath plasma acceleration [Lavraud *et al.*, 2007]. Note that such conditions may lead to both KH and anomalous surface waves. The example of 1/2 magnetosheath speed found in one case of surface waves [Chen *et al.*, 1993] was later applied to other events without distinction. This was done first in another case of anomalous surface wave [Chen and Kivelson, 1993], yielding an estimate of 100 km s^{-1} . This value seems also overestimated given that the observations indicate a time lag of 10 s between ISEE-1 and ISEE-2 (as seen in the work of Chen and Kivelson [1993, Figure 3] that are separated by about 420 km, i.e., a phase speed no greater than about 42 km s^{-1}).

[15] Another ad hoc approximation is to take the average mean bulk flow velocity in the boundary layer as estimate for the phase speed. Several properties of the magnetic and velocity fields show gradually varying profiles across the layer [e.g., Phan *et al.*, 1997; Foullon *et al.*, 2008]. Like the group speed in equation (2), the flow velocity in the boundary layer is taken on average to be intermediate between the magnetosheath and the (stagnant) magnetospheric velocities, despite evidence of plasma acceleration presumably signaling the presence of rolled-up KH waves. For a subset of the 20 November 2001 event [Foullon *et al.*, 2008, interval D], the average flow velocity along the X_{GSM} direction of 212 km s^{-1} [Hasegawa, 2009] is twice as large as the phase speed deprojected in the X_{GSM} direction, i.e., $100 \pm 11 \text{ km s}^{-1}$ ($V_k = 85 \pm 9 \text{ km s}^{-1}$ from Foullon *et al.* [2008]). For coilignment of the time series in the comoving frame of the KH wave in the X_{GSM} direction [Hasegawa *et al.*, 2004], this apparent phase speed discrepancy on the Cluster 2000 km scale size can only lead to a spatial or temporal shift correction between time series of about $(2000 \times (1/100 - 1/212) \sim) 11 \text{ s}$ at most. Therefore this inaccuracy does not noticeably affect the apparent good alignment of magnetic and velocity fluctuations [Hasegawa *et al.*, 2004, Figure 2] with period of 215 s [Foullon *et al.*, 2008]. As a result, it is unlikely to change the observations of higher densities intruding in lower density regions, which is central to the interpretation of rolled-up vortices [Hasegawa *et al.*, 2004]. However, the inaccuracy of the phase speed by a factor greater than two has a direct effect on the wavelength determination.

4. Wavelength Proxy From Two-Point Distant Comparison

[16] Multipoint in situ observations are exceptionally favorable for data analysis and the characterization of KH

activity. However, it becomes more challenging to cross-correlate data sets at increasing separation distances due to the misalignment of the spacecraft along \mathbf{k} . For illustration we use the 20 November 2001 event, with two low-latitude observations separated by several R_E along the dusk flank magnetopause, as shown in Figure 4a, with Geotail on the dayside and Cluster tailward of the terminator. The two-point observations may be compared at the time of relatively large IMF clock angle ($41 \pm 18^\circ$ between 1804 and 2004 UT [Foullon *et al.*, 2008, interval C]).

[17] Figures 2 and 3 illustrate the spectral analysis of the wave activity at Geotail and Cluster, using magnetic field data from the Magnetic Field investigation (MGF) aboard Geotail [Kokubun *et al.*, 1994] and the Fluxgate Magnetometer (FGM) aboard Cluster C2 [Balogh *et al.*, 2001], with time cadence of about 3 s and 0.2 s, respectively. For comparison between the two data sets, the magnetic field vectors \mathbf{B} are first projected into a boundary coordinate system ($\mathbf{l}, \mathbf{m}, \mathbf{n}$): \mathbf{n} is chosen to be the normal to a model magnetopause [Roelof and Sibeck, 1993] scaled to the respective spacecraft position (at 1918 UT, for Geotail, $\mathbf{n} = [0.833, 0.548, 0.075]$ and at 1904 UT, for Cluster, $\mathbf{n} = [0.504, 0.855, -0.120]$, given in the GSM system), $\mathbf{m} = \mathbf{n} \times \mathbf{z}_{GSM}/|\mathbf{n} \times \mathbf{z}_{GSM}|$ and $\mathbf{l} = \mathbf{m} \times \mathbf{n}$ [Elphic and Russell, 1979]. As shown in Figures 2a–2c and Figures 3a–3c, the components are then normalized and interpolated to a regular 1-s cadence time array.

[18] For the whole event at Cluster, we have shown that the magnetic wave activity is characterized with the largest component $B(m)$ of the magnetic field and that amplitude variations in this tailward component (m pointing sunward on the dusk flank) relate to λ changes (due to the tailward component stretching in the wave propagation direction). $B(m)$ is also the component with the largest variations at Geotail, which was in the boundary layer between 1918 and 1955 UT (indicated between vertical dotted lines in Figures 2b and 3b). According to our cross-correlation analysis of the magnetosheath clock angle adjacent to the magnetopause, there is less than 5 min propagation delay along streamlines (adjacent to the magnetopause) from the dayside magnetosheath near Geotail to the nightside magnetosheath near Cluster [Foullon *et al.*, 2008]. Since the wave propagation delay in the boundary layer is expected to be even shorter, the time frame for comparison between the two sets of perturbations is valid.

[19] Prior to spectral analysis, we detrend the time series $B(m)/|B|$ in a wavelet ‘à trous’ decomposition, in order to filter out the low frequencies presumably associated with solar wind-driven variations. The trend is overlaid for each time series in Figures 2b and 3b. Figures 2d and 3d show the Lomb-Scargle periodogram, normalized Fourier power spectrum, and normalized global (time-averaged) Morlet wavelet (GW) spectrum for the detrended time series $\delta(B(m)/|B|)$. In the range 100–250 s (4–10 mHz), we can distinguish power spectral peaks at periodicities near 144.2 s at Geotail and 187 s at Cluster, which are consistent with the sizes of individual oscillatory structures, whose amplitudes dominate in the time series when inspected visually. As noted by Foullon *et al.* [2008], a secondary periodicity ($301 \pm 13 \text{ s}$) is present in the Cluster time series; signals above the periodogram confidence level [Bai and Cliver, 1990] are also present near those periods in the Geotail time series.

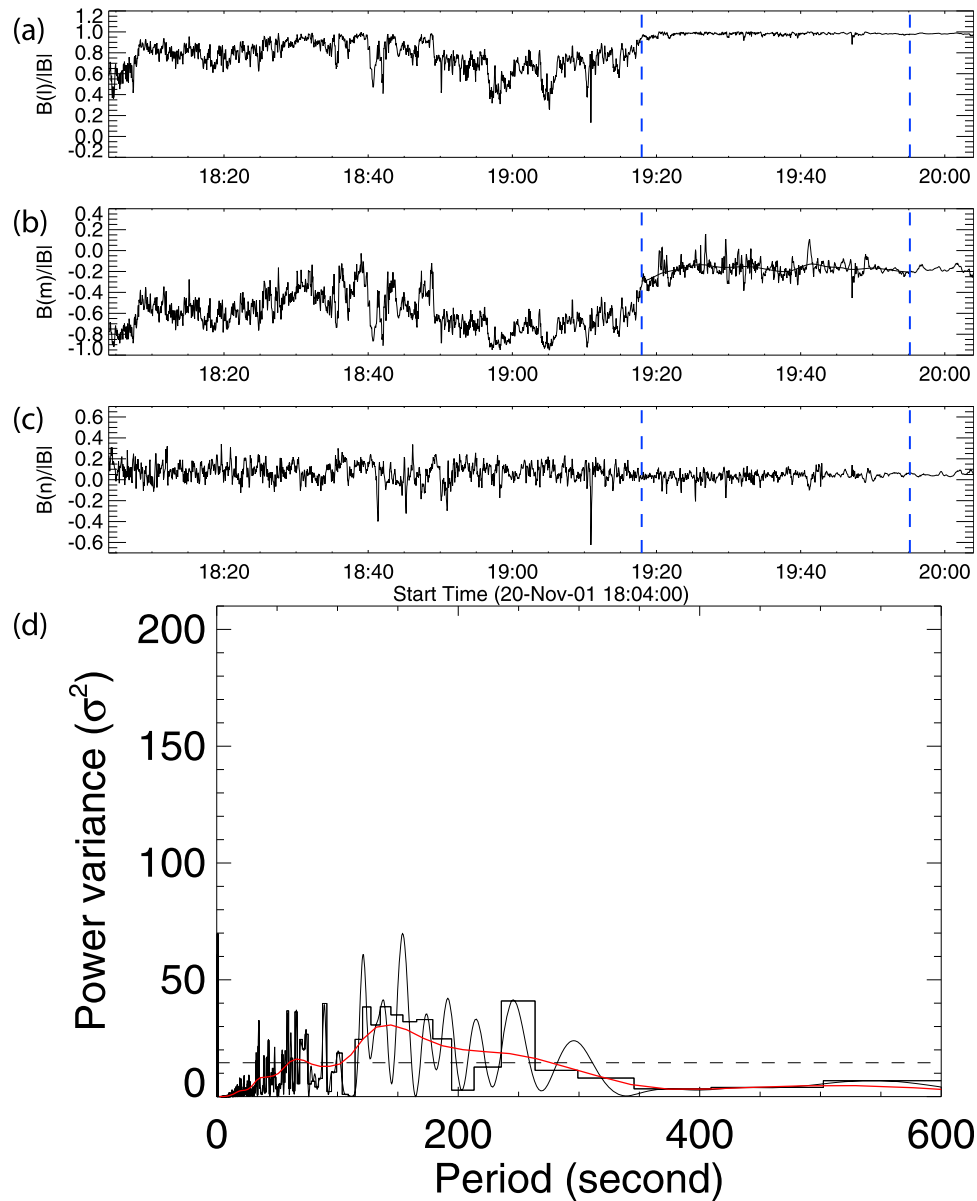


Figure 2. Magnetic wave activity at Geotail in the interval between 1804 and 2004 UT on 20 November 2001. (a–c) Normalized magnetic field time series $B/l|B|$, projected into a lmn coordinate system; a trend is overlaid on the m component, $B(m)/|B|$. Vertical dashed lines mark interval in the boundary layer. (d) Lomb-Scargle periodogram (thin solid), normalized Fourier power spectrum (histogram mode) and normalized global wavelet spectrum (thick red) on the detrended time series $\delta(B(m)/|B|)$. The horizontal dashed line is the 95% confidence level for the periodogram.

[20] Compared to the dominant signals, the secondary signals at Cluster appears with less power in the Fourier and periodogram spectra but with more power in the GW power spectrum. Despite this difference, to keep in mind, the GW spectrum gives a smoothed version of the Fourier power spectrum [Torrence and Compo, 1998]. Thus to facilitate the comparison between Geotail and Cluster spectra, Figure 4b presents the GW spectra together but without normalization. There is a tenfold increase in power at Cluster (interval C) with respect to Geotail (for the interval in the boundary layer). Figure 4b shows also the 43 s

increase in periods from 144.2 s at Geotail to 187 s at Cluster. While the data may be interpreted as a growth of activity from Geotail to Cluster, i.e., the λ of the structures increases, due to an increasing speed of propagation, these distant observations may not be necessarily related to each other. Geotail is near the equator and close to the stagnation region, while the wave activity at Cluster is likely to originate from a position intermediate between Geotail and Cluster (as seen from the subsolar point), since the IMF is substantially away from the Northward direction (see Foullon *et al.* [2008], KH unstable regions).

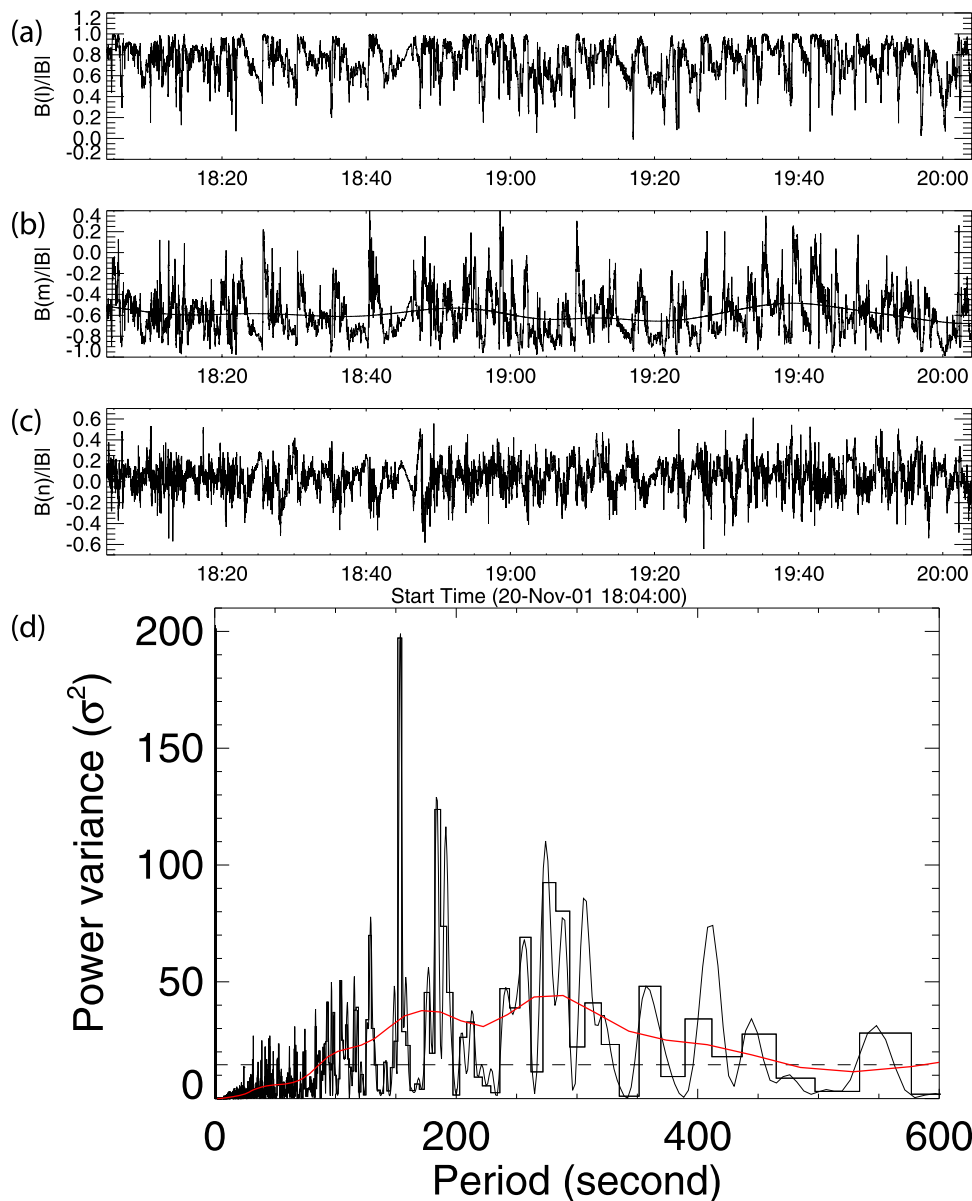


Figure 3. Magnetic wave activity at Cluster C2 in the same interval and same format as in Figure 2.

[21] Nevertheless, using the spectral power as proxy for λ , we can infer the relation between the two parameters at Cluster and apply this relation in order to obtain the λ at Geotail. A statistical fit at Cluster can be obtained thanks to various levels of IMF clock angles, which control λ [Foullon *et al.*, 2008]. Amplitude variations in B_m are assumed to depend on λ , whatever the nature and origin of the perturbations, so that the wave activities from sites, which are not necessarily aligned in a given propagation path, can still be compared. Figure 4c is a plot of the global wavelet maximum power, P_{GW} , in the range 100–250 s (4–10 mHz), versus the wavelengths. The Cluster λ (upper symbols in black) are obtained by combining results from four-spacecraft timing and spectral analysis. The different symbols correspond to the various levels of IMF clock angles, with positive values around (A: diamond) 36°,

(B: square) 12°, (C: triangle) 41°, (D: circle) 19° and (E: asterisk) 62°, corresponding to time intervals with relatively small variations of the solar wind speed and density between them. The best fit to the data points at Cluster, using the linear relation

$$\lambda(P_{GW}) = \lambda_0 + A \log(P_{GW}), \quad (3)$$

with λ in km and P_{GW} in Hz^{-1} , is obtained for $\lambda_0 = 11860 \pm 471$ km and $A = 1739 \pm 487$ km. This relation is used to estimate the λ at Geotail (lower triangle in red), i.e., 7253 ± 690 km. From the dominant spectral period, we infer a propagation speed of the perturbations at Geotail of 50 ± 5 km s^{-1} . This value is slightly lower but close to the average ion velocity projected in the \mathbf{m} direction of the local

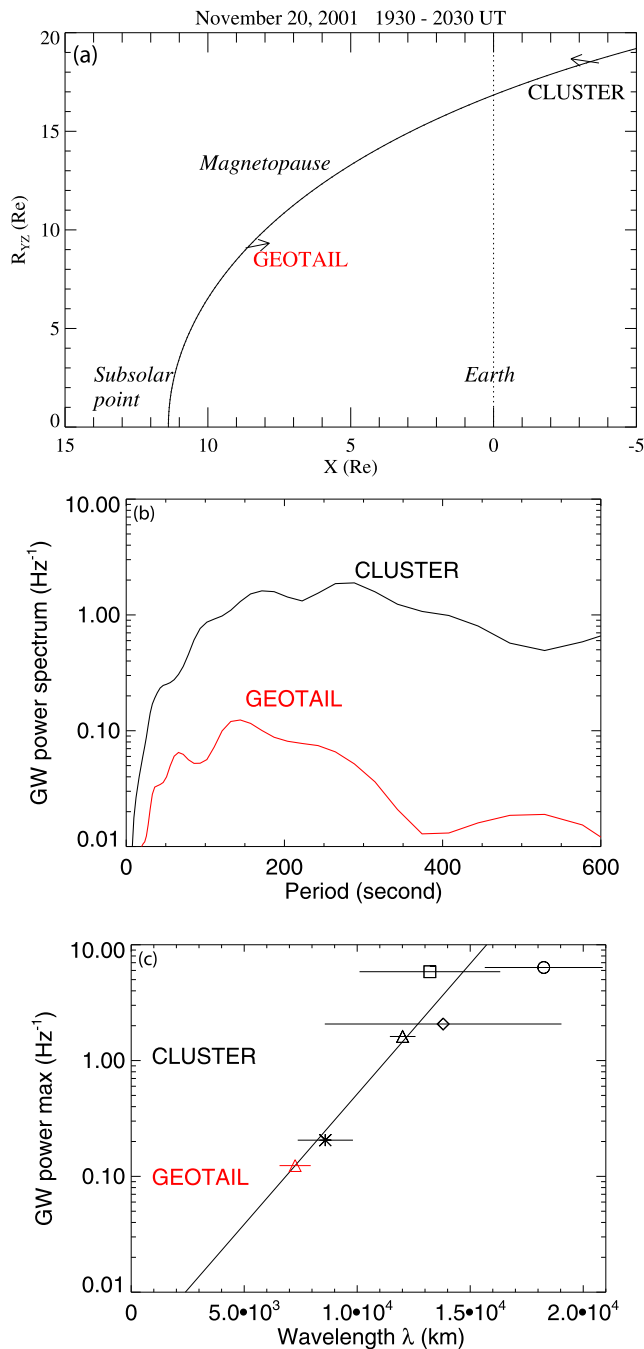


Figure 4. Event of 20 November 2001. (a) Orbit view of Geotail on the dayside and Cluster tailward in a radial cross-section of the dusk magnetosphere. (b) Global wavelet power spectra for the detrended and normalized time series $\delta B(m)/|B|$ at Cluster (black) and Geotail (red). (c) Global wavelet maximum power in the range 100–250 s (4–10 mHz), versus the wavelengths. See text for details.

boundary coordinate system, namely, 57 km s^{-1} in absolute value. Adding error estimates on the spectral power at Geotail (e.g., due to the low-frequency filtering) may slightly broaden the range of λ and the propagation speed of the perturbations. This indicates that the perturbations at Geotail are likely to be advected with the local flow. We note that the value obtained is

much lower than the 115 km s^{-1} estimate given by *Hasegawa et al.* [2009] (based on flow approximations).

5. Conclusions

[22] Although they represent a modest sample of the observations analyzed over the past, the phase speeds derived so far with Cluster yield values of a 1/3 or less of the magnetosheath flow speed or the average flow speed in the boundary layer. Time lags between two or more spacecraft can give a qualitative upper estimate of the phase speed. Cross-examination of a few cases confirm that the flow approximations often used are generally too large. Such inaccuracies may lead to mistakes in the characterization of the different types of waves, in particular KH and anomalous types, and in the sizes of KH rolled-up vortices and other derived quantities of physical importance for the conditions leading to their formation. Using two-point distant observations and spectral analysis of the tailward magnetic field component, we proposed an alternative method to estimate the wavelength and phase speed at a single spacecraft from a statistical fit to the data at the other site. This method might be further validated with more observations.

[23] **Acknowledgments.** C. Foullon acknowledges financial support from the UK Science and Technology Facilities Council (STFC) on the CFSA Rolling Grant. C. J. Farrugia acknowledges financial support from NASA grant NNX08AD11G. We acknowledge use of Geotail and Cluster data, in particular from magnetic field in situ experiments, as shown in this paper, and thank their respective Principal Investigators, namely S. Kokubun (Geotail/MGF) and E. Lucek (Cluster/FGM).

[24] Masaki Fujimoto thanks Johan De Keyser and another reviewer for their assistance in evaluating this paper.

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